MEAN SEPARATION BETWEEN NUCLEATION SITES AND COSMIC QUARK - HADRON PHASE TRANSITION

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Following an analysis of Kajantie and Kurki-Suonio [1] and Alcock et al. [2], the mean separation between the nucleation sites (phases) ℓ , after the nucleation had occured in the expansion of the Universe, is given by [2]

$$\ell = 0.3 \sigma^{3/2} t / T_c^{1/2} L,$$
 (1)

where t is the age of the Universe, T_C is the phase transition temperature, L is the latent heat, and σ is the surface tension. This length ℓ will assume a greater significance in the calculations of the light element abundances of the early Universe using non-standard big bang theory [1-4].

al. [3] have calculated the amount of Fuller et latent heat released, during the guark - hadron transition, for different phase coexistence temperatures T_c. In their calculations, the latent heat L is fixed for a given T_c (see Table I of Ref. To obtain the appropriate value of nucleation 3). sites for PNS calculations, they have to change the of T_c. But in our lowest order perturbative value QCD model [5], we can calculate different values of for a given T_C by adjusting the QCD scale fixing parameter Λ . Our calculated values of T_c, and L phase transition time Δt for different values of В and Λ are presented in Table I. The calculated values of latent heat released during the phase transition is always larger than that of Ref.6, this being due to the inclusion of the proper volume correction (Hagedorn) in the HRG phase. Addition more and more hadronic resonances in the hadronic phase decreases the release of latent heat during the cosmic quark-hadron phase transition. Inclusion

of Hagedorn's correction, however, decreases the critical energy density. Our calculation presented in Table I may be useful to estimate the appropriate distance between nucleation sites [1-4] and to reproduce the correct values of the light element abundances using the non-standard big bang model.

| В | Λ | T _C | L | Δt |
|--------------|-------|----------------|------------------------|------|
| (MeV/fm^3) | (MeV) | (MeV) | (MeV^4) | (us) |
| 30 | 100 | 102 | 8.95×10^8 | 93 |
| 60 | 0 | 104 | 1.81 × 10 ⁹ | 98 |
| 60 | 250 | 150 | 1.59×10^9 | 23 |
| 150 | 100 | 150 | 4.30 x 10^9 | 29 |
| 250 | 0 | 150 | 7.30×10^9 | 97 |

References:

Table T.

| 1. | K. Kajantie and H. Kuriki-Suonio, Phys. Rev. D34, |
|----|---|
| | 1719 (1986). |
| 2. | C. Alcock et al., Nucl. Phys. A498, 301 (1989). |
| 3. | G. M. Fuller et al., Phys. Rev. D37, 1380 (1988). |
| 4. | T. Kajino, Phys. Rev. Lett. 66, 125 (1991). |
| 5. | K. Sakthi Murugesan, G. Janahvi, and |
| | P. R. Subramanian, Phys. Rev. D41, 2384 (1990); |
| | D42, 3576 (1990). |

6. H. Voss et al., J. Phys. G15, 561 (1989).